

The Slow Winds of A-Type Supergiants

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Abstract. The line driven- and rotation modulated-wind theory predicts an alternative *slow* solution, besides from the standard m-CAK solution, when the rotational velocity is close to the critical velocity. We study the behaviour of the winds of A-type supergiants (Asg) and show that under particular conditions, e.g., when the δ line-force parameter is about 0.25, the slow solution could exist over the whole star, even for the cases when the rotational speed is slow or zero. We discuss density and velocity profiles as well as possible observational counterparts.

Keywords. stars: A supergiants, winds, mass-loss rates

1. Introduction

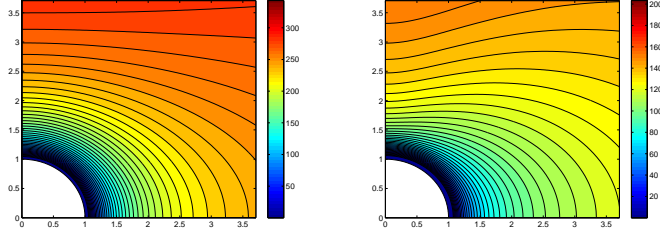
The theory of radiation driven winds (CAK, Castor et al. 1975) and later m-CAK (Pauldrach et al. 1986) succeeded in describing terminal velocities (V_∞) and mass losses (\dot{M}) of hot stars apart from predicting the wind momentum luminosity (WML) relationship. This relationship was empirically found for the most of luminous O-type stars (Puls et al. 1996) and extended to lower luminosity objects by Kudritzki et al. (1999), who found that it depends on the spectral type. Particularly, Asg show V_∞ values a factor 3 lower than theoretical ones (Achmad et al. 1997) and V_∞ decreases when increasing the escape velocity (V_{esc}) (Verdugo et al. 1999), in clear contradiction with the CAK theory (fast solution). Moreover, Kudritzki et al. (1999) found that $H\alpha$ profile of these stars can be modeled with β velocity laws, for $\beta > 1$. These observational discrepancies with CAK and m-CAK theories could be related to a change in the parameter α along the wind due either to a change in the ionization of the wind or to a decoupling of the line-driven ions in the wind from the ambient gas (Achmad et al. 1997). In 2004, Curé (2004) revisited the theory of steady fast-rotating line-driven winds and found that for $\omega = V/V_{crit} > 70\%$ there exists another hydrodynamical or *slow* solution, which is denser and slower than the standard m-CAK solution. For B-type stars he obtained velocity distributions, as well as the critical point and \dot{M} , that can be matched by a velocity law with $\beta > 1$. Therefore, we propose here that the winds of Asg could be related to the *slow hydrodynamical solutions* described by Curé (2004). We demonstrate numerically that for Asg, given a particular set of wind parameters, k , α and δ , there exists a *slow solution* at all latitudes for low rotational velocities ($\omega < 0.4$), and even for the non-rotating case.

2. Results and Conclusions

We calculated numerically wind solutions for different rotational velocities and different stellar and wind parameters. Table 1 lists some of the models we computed for low values of ω : those marked with an "s" in the first column present a slow solution at all latitudes, while the remaining corresponds to the fast solution at all latitudes ("f"). Figure 1 displays the behaviour of the fast and slow solutions as a function of latitude for models with

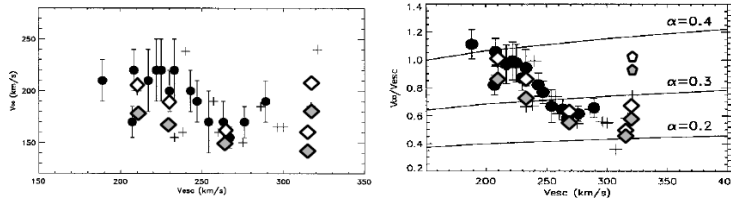
Table 1. Stellar and Wind Parameters

Mod	Teff [K]	log g	R [R _⊙]	ω	α	k	δ	α_{eff}	$V_{\infty pol}$ [km s ⁻¹]	$F_{m, pol}$ [M _⊙ /(yr sr)]	$V_{\infty eq}$ [km s ⁻¹]	$F_{m, eq}$ [M _⊙ /(yr sr)]	\dot{M} [M _⊙ /yr]
1 (s)	13000	1.73	68	0.4	0.51	0.03	0.23	0.28	160	4.78×10^{-5}	142	5.78×10^{-5}	6.41×10^{-4}
2 (f)	10000	2.0	60	0.4	0.49	0.07	0.15	0.34	350	1.14×10^{-5}	291	1.54×10^{-5}	1.59×10^{-4}
3 (s)	10000	2.0	60	0.4	0.49	0.07	0.26	0.23	207	9.79×10^{-8}	181	1.21×10^{-7}	1.32×10^{-6}
4 (s)	9500	1.7	80	0.4	0.49	0.07	0.26	0.23	168	5.94×10^{-7}	149	7.40×10^{-7}	8.04×10^{-6}
5 (s)	9000	1.7	100	0.4	0.49	0.07	0.26	0.23	188	3.08×10^{-7}	165	3.84×10^{-7}	4.18×10^{-6}
6 (s)	9000	1.7	120	0.4	0.49	0.07	0.26	0.23	206	3.57×10^{-7}	178	4.49×10^{-7}	4.87×10^{-6}

**Figure 1.** Wind velocity distributions. Left: Model 2 (Fast). Right: Model 3 (Slow).

the same stellar and wind parameters, except δ . \dot{M} and V_{∞} obtained with models 1, 3, 4, 5 and 6 are in good agreement with those observed in Asg (Verdugo et al. 1999; Kudritzki et al. 1999). Instead, the fast solution leads to higher \dot{M} and V_{∞} . We find that the solutions for slow winds (see Figure 2, in white/grey romboids) match observational data of some Asg stars (black symbols, Verdugo et al. 1998).

By solving the 1-D hydrodynamic equation of rotating line-driven winds of Asg for different sets of line force parameters, we found a slow wind regime over all latitudes when increasing the parameter δ , which describes changes in the ionization stage of the wind. These slow solutions trace the observational trends found by Verdugo et al. (1999). Our result supports these authors' hypothesis, stating that the negative slope of V_{∞}/V_{esc} vs. V_{esc} could be related to the degree of ionization and wind density.

**Figure 2.** The white/grey romboids (pentagons) correspond to equator/pole slow (fast) solutions, black symbols correspond to Figs. 2 and 3 of Verdugo et al. (1998)

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